- *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)
- Under presence of *M*, Newton's laws of the forces on *m*

프 🖌 🗶 프 🛌

æ

• *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)

 Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1):

э.

• *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)

• Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1): $Mm - \vec{X} \qquad d^2\vec{X}$

$$\frac{1}{r^2} \frac{-x}{r} = m \frac{d^2 x}{dt^2}$$
where $r = \sqrt{x^2 + v^2 + z^2}$

▶ < 프 > < 프 > · · 프

- *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0,0,0)
- Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1):

$$\frac{Mm}{r^2}\frac{-\dot{X}}{r} = m\frac{d^2\dot{X}}{dt^2}$$

where
$$r = \sqrt{x^2 + y^2 + z^2}$$

• Define $\varphi(r) = -\frac{M}{r}$ as the potential function We'll prove Laplace's equation: $\nabla^2 \varphi = \frac{\partial^2 \varphi}{\partial x^2} + \frac{\partial^2 \varphi}{\partial y^2} + \frac{\partial^2 \varphi}{\partial z^2} =$

- *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)
- Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1):

$$\frac{Mm}{r^2}\frac{-\dot{X}}{r} = m\frac{d^2\dot{X}}{dt^2}$$

where
$$r = \sqrt{x^2 + y^2 + z^2}$$

• Define $\varphi(r) = -\frac{M}{r}$ as the potential function We'll prove Laplace's equation: $\nabla^2 \varphi = \frac{\partial^2 \varphi}{\partial x^2} + \frac{\partial^2 \varphi}{\partial y^2} + \frac{\partial^2 \varphi}{\partial z^2} = 0$ Compute $\frac{\partial \varphi}{\partial x}$ by chain rule $\frac{\partial \varphi}{\partial r} \frac{\partial r}{\partial x}$ Then compute $\frac{\partial^2 \varphi}{\partial x^2}$ And similar for y and z. Add.

- *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)
- Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1):

$$\frac{Mm}{r^2}\frac{\dot{-X}}{r} = m\frac{d^2\dot{X}}{dt^2}$$

where
$$r = \sqrt{x^2 + y^2 + z^2}$$

- Define $\varphi(r) = -\frac{M}{r}$ as the potential function We'll prove Laplace's equation: $\nabla^2 \varphi = \frac{\partial^2 \varphi}{\partial x^2} + \frac{\partial^2 \varphi}{\partial y^2} + \frac{\partial^2 \varphi}{\partial z^2} = 0$ Compute $\frac{\partial \varphi}{\partial x}$ by chain rule $\frac{\partial \varphi}{\partial r} \frac{\partial r}{\partial x}$ Then compute $\frac{\partial^2 \varphi}{\partial x^2}$ And similar for y and z. Add.
- Corollary: Consider $-\nabla \varphi$ and its relationship to $\frac{d^2 X}{dt^2}$

通 と く ヨ と く ヨ と

- *m* at $\vec{X}(t) = (x(t), y(t), z(t))$ and *M* at (0, 0, 0)
- Under presence of *M*, Newton's laws of the forces on *m* (Einstein set units so G=1):

$$\frac{Mm}{r^2} - \frac{X}{r} = m \frac{d^2 X}{dt^2}$$

where
$$r = \sqrt{x^2 + y^2 + z^2}$$

- Define $\varphi(r) = -\frac{M}{r}$ as the potential function We'll prove Laplace's equation: $\nabla^2 \varphi = \frac{\partial^2 \varphi}{\partial x^2} + \frac{\partial^2 \varphi}{\partial y^2} + \frac{\partial^2 \varphi}{\partial z^2} = 0$ Compute $\frac{\partial \varphi}{\partial x}$ by chain rule $\frac{\partial \varphi}{\partial r} \frac{\partial r}{\partial x}$ Then compute $\frac{\partial^2 \varphi}{\partial x^2}$ And similar for y and z. Add.
- Corollary: Consider $-\nabla \varphi$ and its relationship to $\frac{d^2 \vec{X}}{dt^2}$ $\frac{d^2 x}{dt^2} = -\frac{\partial \varphi}{\partial x}, \frac{d^2 y}{dt^2} = -\frac{\partial \varphi}{\partial y}$, and $\frac{d^2 z}{dt^2} = -\frac{\partial \varphi}{\partial z}$

- Finite # point masses: φ is a sum in Laplace's eq
- Einstein replaced the corollary with

$$rac{d^2 x^\lambda}{ds^2} = -\Gamma^\lambda_{\mu
u} rac{dx^\mu}{ds} rac{dx^
u}{ds}$$

同 ト イヨ ト イヨ ト ヨ うくで

- Finite # point masses: φ is a sum in Laplace's eq
- Einstein replaced the corollary with

$$\frac{d^2 x^{\lambda}}{ds^2} = -\Gamma^{\lambda}_{\mu\nu} \frac{dx^{\mu}}{ds} \frac{dx^{\nu}}{ds}$$

- $\frac{\partial \varphi}{\partial x^i} \& \Gamma^{\lambda}_{\mu\nu} \frac{dx^{\mu}}{ds}$ similar roles: Field equations relate these potential functions to the distribution of matter
- 2nd order PDE 16 eqs and 16 unknowns—solve for metric

- Finite # point masses: φ is a sum in Laplace's eq
- Einstein replaced the corollary with

$$rac{d^2 x^\lambda}{ds^2} = -\Gamma^\lambda_{\mu
u} rac{dx^\mu}{ds} rac{dx^
u}{ds}$$

- $\frac{\partial \varphi}{\partial x^i} \& \Gamma^{\lambda}_{\mu\nu} \frac{dx^{\mu}}{ds}$ similar roles: Field equations relate these potential functions to the distribution of matter
- 2nd order PDE 16 eqs and 16 unknowns—solve for metric
- If $\Gamma^{\lambda}_{\mu\nu} = 0$ then $\frac{d^2x^{\lambda}}{ds^2} = 0$
- Field equations written in the Christoffel symbols:

$$\frac{\partial \Gamma^{\lambda}_{\mu\lambda}}{\partial x^{\nu}} - \frac{\partial \Gamma^{\lambda}_{\mu\nu}}{\partial x^{\lambda}} + \Gamma^{\beta}_{\mu\lambda}\Gamma^{\lambda}_{\nu\beta} - \Gamma^{\beta}_{\mu\nu}\Gamma^{\lambda}_{\beta\lambda} = 0$$

御入 くさん くさん 一注

Consequences and Experiments

- Spacetime in the presence of masses is curved and geodesics more interesting
- Gravity is the curvature of spacetime

프 에 에 프 어 - -

ъ

Consequences and Experiments

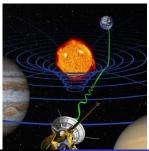
- Spacetime in the presence of masses is curved and geodesics more interesting
- Gravity is the curvature of spacetime
- Arthur Eddington (1919): star near sun shifted by amount predicted by relativity! → Einstein public figure
- Radio sources

個 とく ヨ とく ヨ とう

1

Consequences and Experiments

- Spacetime in the presence of masses is curved and geodesics more interesting
- Gravity is the curvature of spacetime
- Arthur Eddington (1919): star near sun shifted by amount predicted by relativity! → Einstein public figure
- Radio sources
- Gravitational lensing and LIGO gravitational waves
- Precession of the orbit of Mercury



Dr. Sarah Math 4140/5530: Differential Geometry

How many general-relativity theorists does it take to change a light bulb?



Dr. Sarah Math 4140/5530: Differential Geometry

(E) < (E)</p>

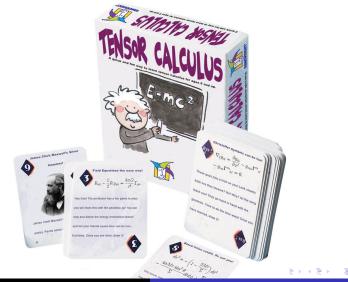
ъ

How many general-relativity theorists does it take to change a light bulb?



Two. One to hold the bulb and one to rotate space.

 Einstein's general relativity remains scientists' best understanding of gravity and a key to our understanding of the cosmos on the grandest scale [Theory of Everything?]



Course Overview

- Curves: torsion and curvature, piece together (DEs) to give us a nice curve. Frenet frame basis
- Surfaces: 1st and 2nd fundamental forms. 1st and Christoffel symbols intrinsic → Gauss curvature. *E*, *F*, *G* piece together to give us a nice smooth surface, or one with manageable singularities like the cone. *I*, *m*, *n* determine how the surface sits in space (if it does)
- SpaceTimes: g_{ij} and Γ^k_{ij} more terms → curvature tensors. Einstein's field equations: way to solve for "nice" solutions.



Final Research Pres: review + extension + bib + peer review

Dr. Sarah Math 4140/5530: Differential Geometry